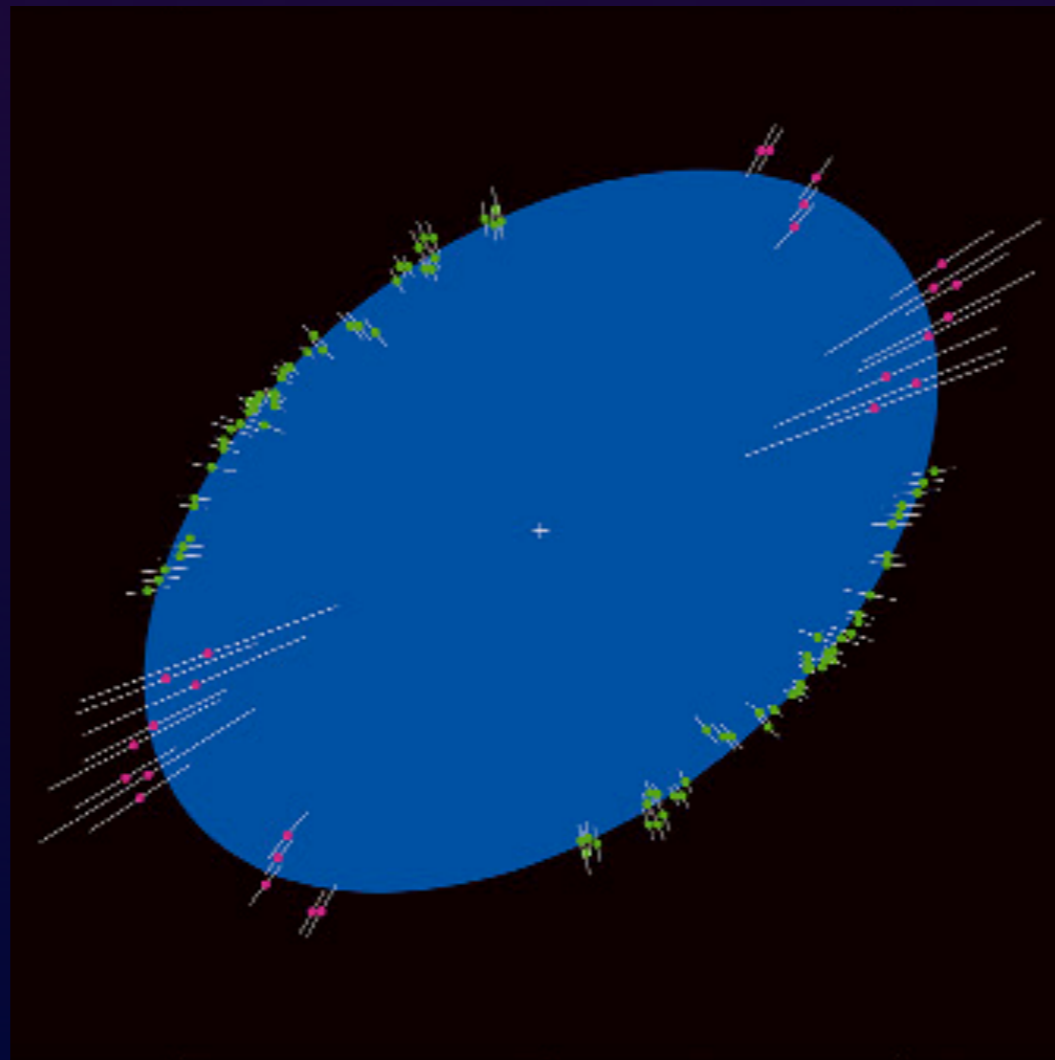
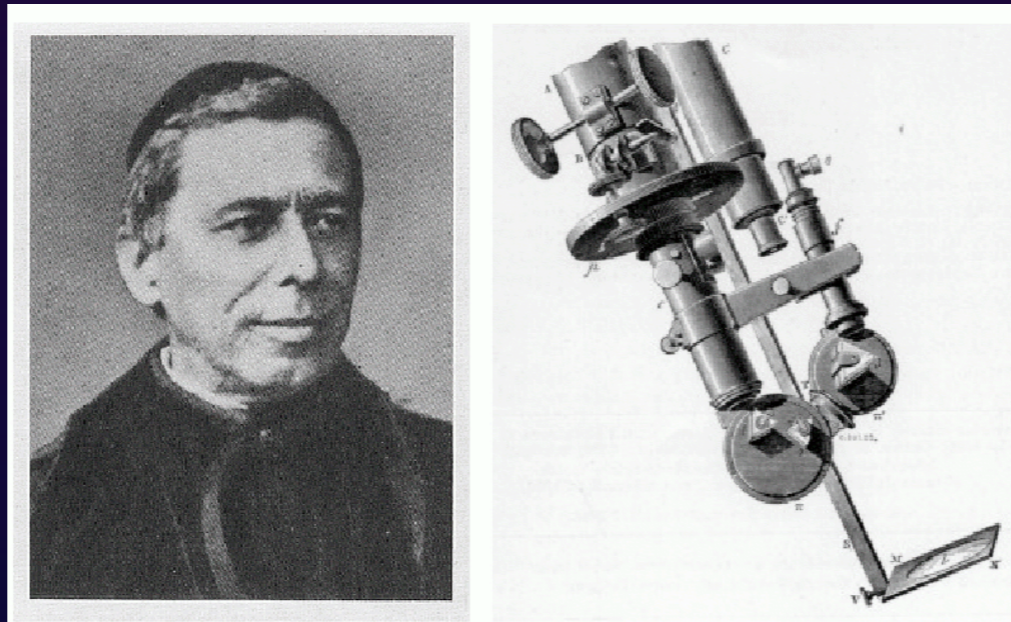


Be stars



Marlene Obbrugger

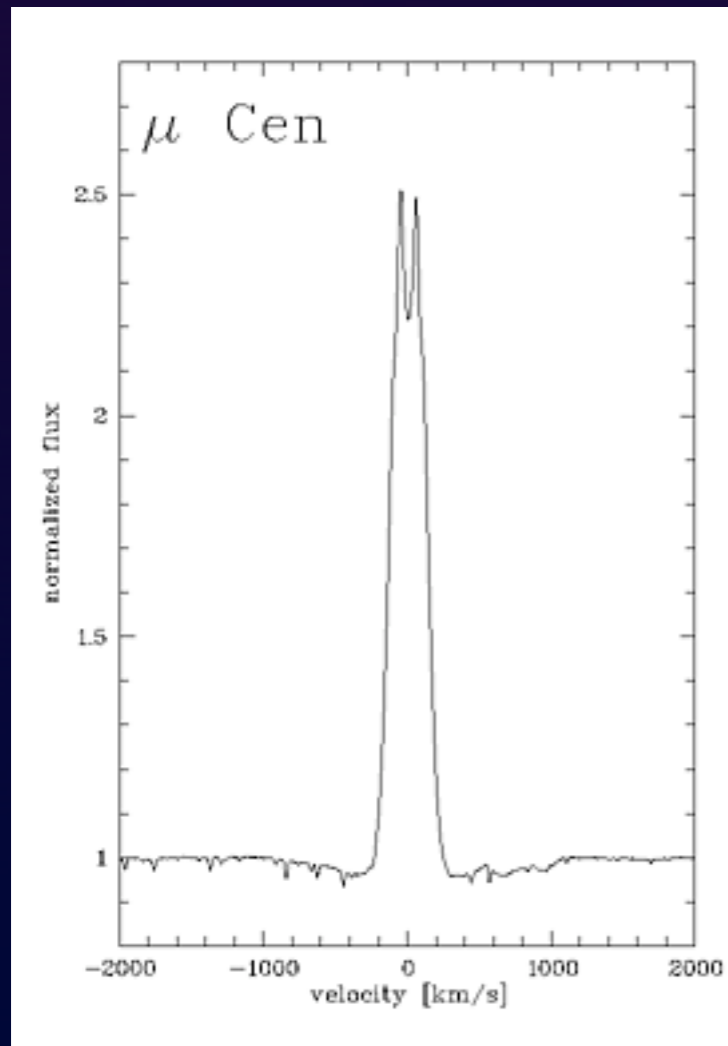
history



Angelo Secchi 1866

first report of a Be star (γ Cas)
visual inspection through a spectroscope

definition of a classical Be star



- at least once Balmer line emission displayed
- non-supergiant B stars

corresponds to about 20 % of all B stars

definition of the B[e] phenomenon

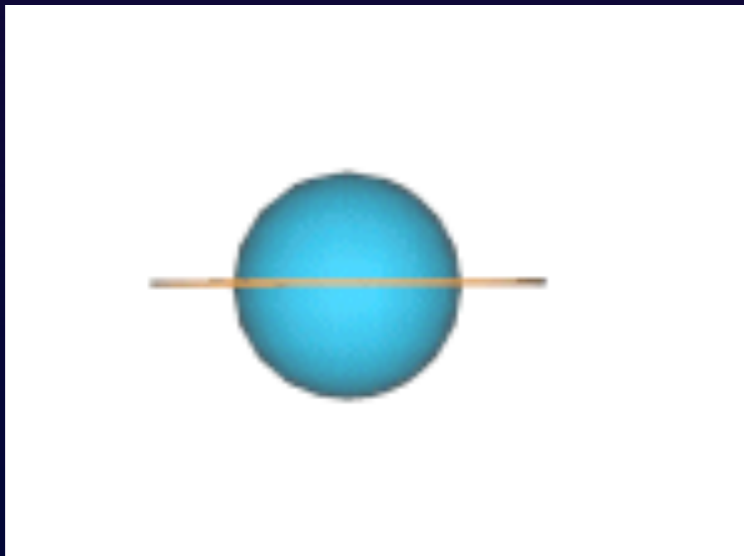
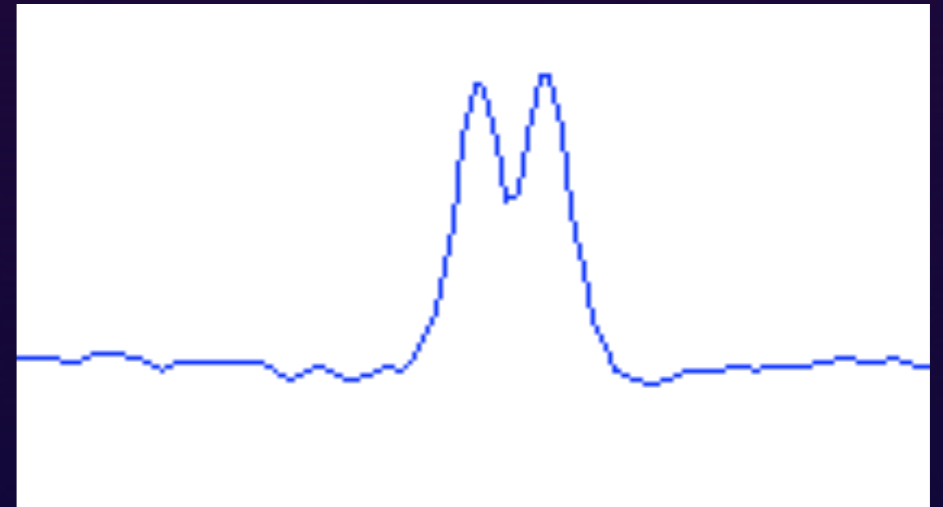
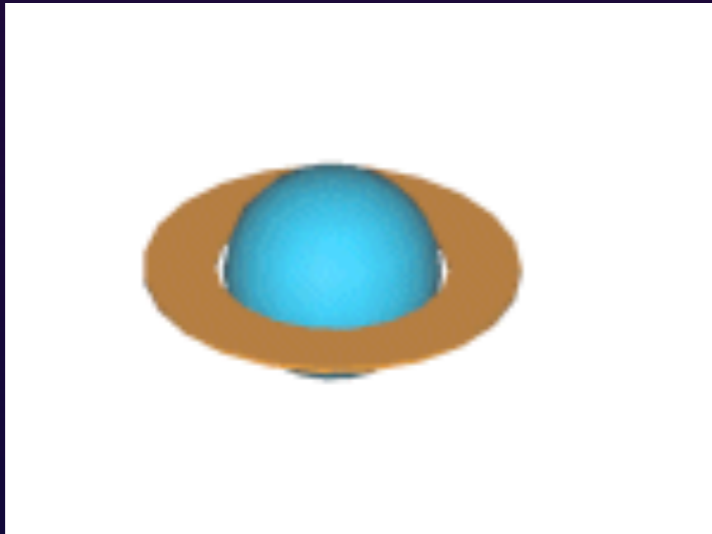
Balmer line emission
forbidden emission lines [FeII] [OI]

applies to: B[e] supergiants
Herbig Ae/Be stars
planetary nebulae
symbiotic stars

properties

- extremely high rotational velocity ($\sim 80\% v_{bu}$)
- emission features due to **circumstellar disk**
- go through different phases:
B-normal \rightarrow Be \rightarrow Be-shell \rightarrow B-shell
- variable spectroscopic, spectrophotometric, polarimetric and interferometric characteristics

shell feature



feature depends on the position of the observer

disk formation theories

radiatively driven winds (not sufficient)

photospheric pulsation reduces gravity locally
(+ driven by radiation)

matter transported along magnetic field lines ?

mass loss due to binary

pulsation → turbulence → angular momentum transfer

MWC297 model (slice through an edge-on view)

Star :

B1.5IVe 23700K
6R_o 10M_o
Rotational velocity = 400km/s

Accretion disc

Inner radius = 16 R_{*}
Outer radius = 50 AU
Opening angle = 5°
Accretion rate : $2 \cdot 10^{-5}$ M_o/year

Equatorial region (disc/wind interface):
Terminal velocity = 70km/s
Mass flux = $0.8 \cdot 10^{-9}$ M_o/year

Stellar Wind :

Mass loss : $3 \cdot 10^{-8}$ M_o/year

Polar region :
Terminal velocity = 600km/s
Mass flux = $3 \cdot 10^{-9}$ M_o/year

evolutionary state

seem to be quite homogeneously distributed along the main sequence (but debate goes on)

late-type Be stars need higher v/v_{crit} and larger t/t_{ms} ratio

massive Be stars tend to be present at lower t/t_{ms} age ratios

variations

1. formation and dispersal of the disk (weeks - decades)
2. binary motion of some stars (days - weeks)
3. non-radial pulsation (0.3 - 2 days)
4. rotation (0.3 - 2 days)

rotation and binarity

- Be stars have higher v_{rot} than B stars
- 75% might have been spun-up by binary mass transfer
- Be star binaries do not exceed 1/3 (typical for non Be)
- rest supposedly rapid rotating at birth

open questions

Does Be frequency rise toward the TAMS ?

Rotational velocity higher than 70%-80% of v_{bu} ?

What fraction has been spun up by binaries ?

Presence and origin of magnetic fields ?

Firm theoretical basis for mass loss due to magnetic field ?



Thank you !

Credit: A. Meilland

The Be Star Alpha Arae
(Artist's Impression)